

EXPERIMENTAL INVESTIGATION OF POTENTIAL HCLAND HF MINERAL BUFFERS ON VENUS

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S (at %) | F (at %) | Cl (at %) |

nd

nd

0.1

S (at %) | F (at %) | Cl (at %) |

nd

nd

0.1

nd

S (at %) F (at %) Cl (at%)

nd

1.9^{SO4}

 $12HCl+5Na_{4}[AlSiO_{4}]_{3}Cl+6CaSiO_{3}=17NaCl+6CaAl_{2}Si_{2}O_{8}+3NaAlSi_{3}O_{8}+6H_{2}O_{1}$

HCl + sodalite + wollastonite = halite + anorthite + albite + water

nd

nd

2.9^{SO4}

1.9^{SO4}

3.1^{SO4}

2.2^{SO4}

1.9^{SO4}

Table 3: Wollastonite was selected as one of the

minerals to address Objective 1. The above

tables show XPS data before (A) and after (B)

the experiment. S is present as a sulfate on both

the chip and powder after the experiment. Cl is

not present on any sample. Some F is present at

Table 4: The sodalite + wollastonite mixture was

selected as one of the mixtures to address

Objective 3 and has been theorized to react via

the equation below. The above table shows XPS

data before and after the experiment. S is present

as a sulfate after the experiment. F is not detected

in the sample. The Cl abundance after the test is

similar to the abundance before the test

depth in the chip, but not in the powder

Figure 1: Wollastonite before (left) and after

(right) the experiment

XPS

Wollastonite Chip

sputtered 1 min (100 Å)

Wollastonite Chip

60 Days in GEER

sputtered 1 min (100 Å)

XPS

Wollastonite Powder

sputtered 1 min (100 Å)

Wollastonite Powder

60 Days in GEER

sputtered 1 min (100 Å)

sputtered 2 min (200 Å)

XPS

Sodalite + Wollastonite

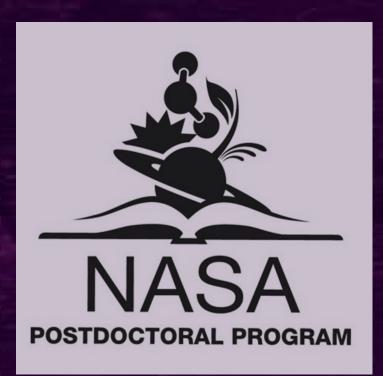
Powder

sputtered 1 min (100 Å)

Sodalite + Wollastonite

Powder 60 days in GEER

sputtered 1 min (100 Å)



Introduction

- \clubsuit The abundance of HCl in Venus' atmosphere is 0.4 \pm 0.03 ppmv and HF is 5 ± 3 ppbv
- * Measurements of HCl and HF in the atmosphere of Venus taken over 20 years apart reveal similar abundances suggesting that both gases are in equilibrium
 - * This alludes to a yet to be determined chlorine and fluorine buffer system with the surface
- Several researchers have used thermodynamic modeling to investigate the interactions between these gases and potential minerals on the surface of Venus [1, 3-7]
 - * However, thermodynamic calculations experimental data to model reactions, and very little experiments have been completed at Venus relevant conditions

Objectives:

- 1. Investigate any interactions that may occur between a single mineral with either HCl or HF to produce a chlorine- or fluorine-bearing mineral
- 2. Investigate if the hydroxyl group in some hydrated silicates can interact with HF and be replaced with fluorine
- 3. Investigate if mineral mixtures can be sinks for Clminerals or F-minerals

Methods

Samples:

- **❖** Total: 27
- Mineral chips were cut and polished using SiC, diamond paste, and colloidal silica to a 50 nm polish
- * Powders were ground using a mortar & pestle and sieved to 100 µm

Glenn Extreme Environment Rig (GEER) [8-9]:

- ❖ Temperature and Pressure: 460°C and ~93 bar
- ❖ Gas: CO₂, N₂, SO₂, H₂O, CO, COS, H₂S, HCl, and HF
- Length of Time: 60 days
- Gas Chromatograph (GC)

Instruments for Analysis:

- ❖ X-Ray Diffraction (XRD)
- Scanning Electron Microscope (SEM)
- Energy-Dispersive x-ray Spectroscopy (EDS)
- ❖ X-ray Photoelectron Spectroscopy (XPS)
- * Brunauer, Emmett and Teller (BET) surface area analysis

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ı	Objective 1	Objective 2	Objective 3
ı	Nepheline	Ferrohornblende	Wollastonite + Sodalite
ı	NaAlSiO ₄	$Ca_{2}(Fe_{4}^{2+}Al)(Si_{7}Al)O_{22}(OH)_{2}$	$CaSiO_3 + Na_4[AlSiO_4]_3Cl$
ı	Albite	Tremolite	Microcline + Enstatite
ı	NaAlSi ₃ O ₈	$Ca_2Mg_5Si_8O_{22}(OH)_2$	$KAlSi_3O_8 + MgSiO_3$
ı	Enstatite	Phlogopite	Microcline + Forsterite
ı	$MgSiO_3$	$KMg_3AlSi_3O_{10}(OH)_2$	$KAlSi_3O_8 + Mg_2SiO_4$
ı			Nepheline + Diopside +
ı	Wollastonite	Pargasite	Enstatite
ı	CaSiO ₃	$NaCa_2Mg_4Al_3Si_6O_{22}(OH)_2$	$NaAlSiO_4 + CaMgSi_2O_6$
ı			+MgSiO ₄
ı	Forsterite	Muscovite	Quartz + Diopside + Enstatite
ı	Mg_2SiO_4	$KAl_3Si_3O_{10}(OH)_2$	$SiO_2 + CaMgSi_2O_6 + MgSiO_9$
		Brucite	
		Mg(OH)	

Table 1: All samples that were selected to address Objectives 1, 2, and 3 in this project

CO_2	N_2	SO ₂	CO	COS	H ₂ S	H ₂ O	HCl	HF
96.5	3.5 ±	150 ±	17 ±	4.4 ±	3 ±	30	0.4	5
±	0.8 %	30	1.4	1	2	ppmv	ppmv	ppbv
0.8%		ppmv	ppmv	ppmv	ppmv			

Table 2: Starting gas composition in the 60-day GEER test

Results

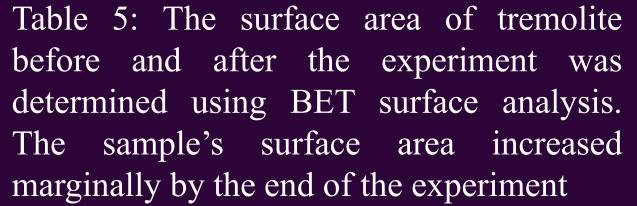
We are unable to display all of the data due to the number of samples investigated in this project. We have chosen a few samples that address each of the objectives to present on this poster. All data is preliminary and further analysis is currently underway

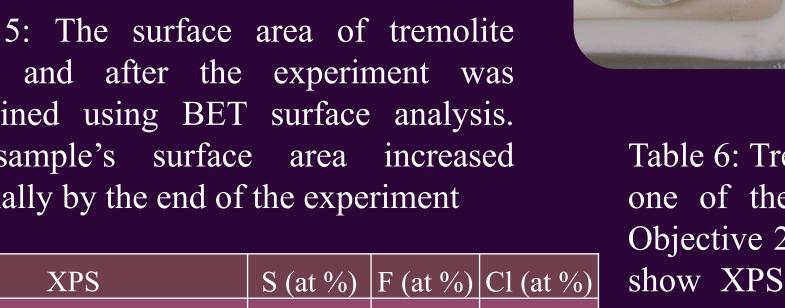
Tremolite Chip



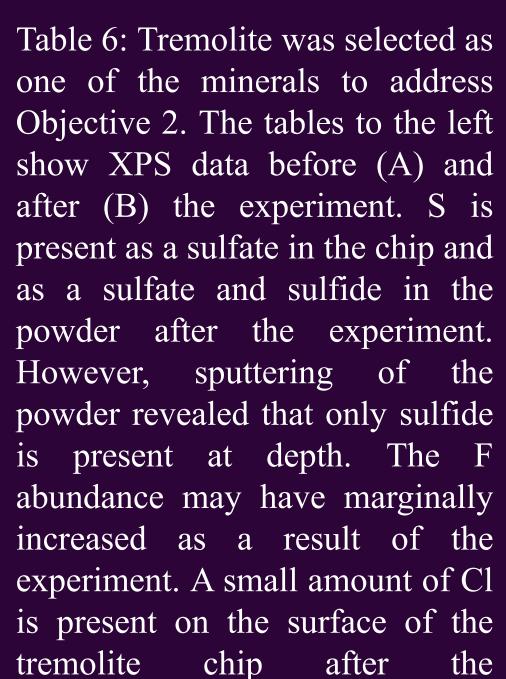
Figure 2: Tremolite before (left) and after (right) the experiment

XPS	Surface Area (m²/g)
Tremolite Powder	$5.0888 \text{ m}^2/\text{g}$
Tremolite Powder	
60 days in GEER	$5.5730 \text{ m}^2/\text{g}$





	110	0.5	110
sputtered 1 min (100 Å)	nd	0.3	nd
Tremolite Chip			
60 days in GEER	0.8SO4	0.8	0.1
sputtered 1 min (100Å)	<.1 ^{SO4}	0.6	<.1
B XPS	S (at %)	F (at %)	Cl (at %)
Tremolite powder	nd	0.2	nd
			116
sputtered 1 min (100 Å)	nd	0.2	nd
sputtered 1 min (100 Å) Tremolite powder	nd	0.2	
	nd 1 ^{SO4/S} -	0.2	



experiment

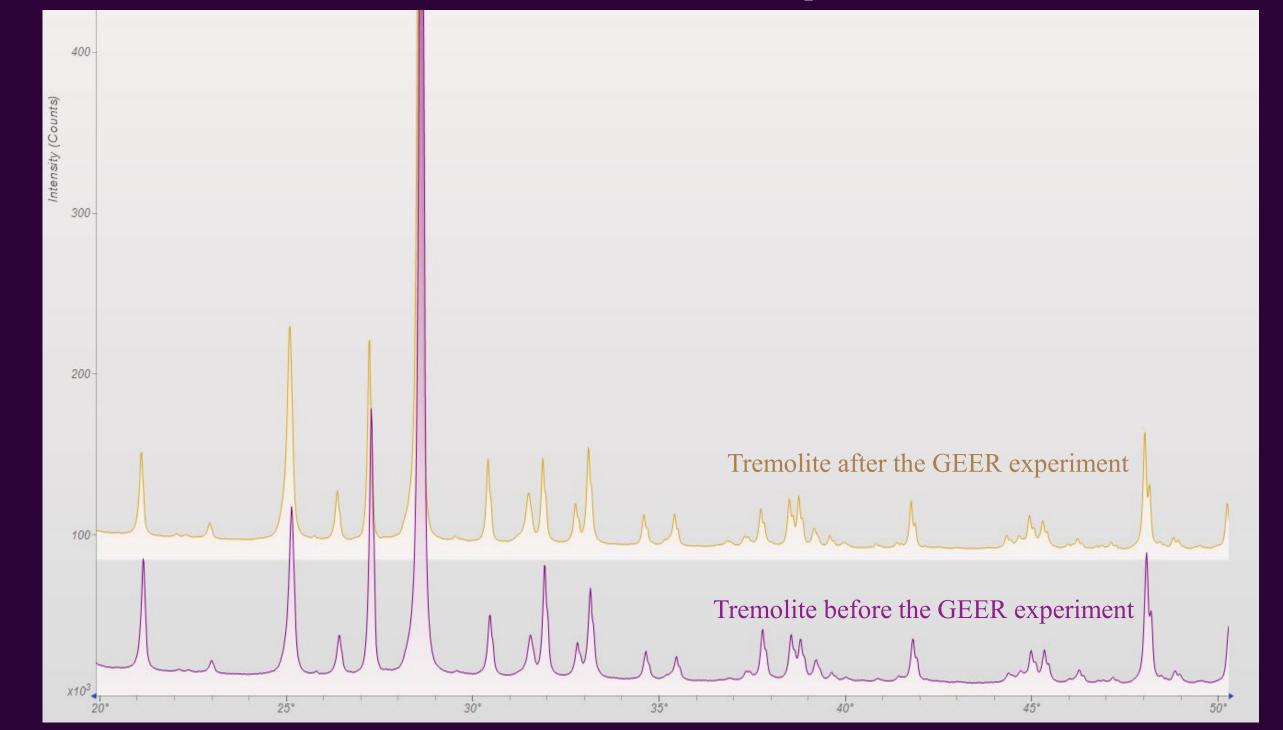


Figure 3: Zoomed in XRD spectra of tremolite before and after the GEER test. No differences were observed between the two spectra. If any changes did occur as a result of the experiment, it was not significant enough to be detected by XRD

Conclusions

- * The vast majority of the samples exhibited a color change into a dull white/grey after the experiment
- * At this time, it cannot be determined if any of the fluorine or chlorine are merely adhered to the surface of the samples, or interacted with the samples to produce secondary minerals
- Sulfur is present on all samples, but the oxidation state varies depending on the cations present in the sample
- Further analysis using SEM/EDS, microprobe and nanoSIMs are planned in the future

References

[1] Fegley Jr., B., Treiman, A. H. (1992) Venus and Mars: Atmosphere, Ionospheres, and Solar Wind Interactions, Geophysical Monograph, 66, 7-72. [3] Barsukov, V. L. et al., (1982) J. Geophys. Res. Solid Earth 87, A3-A9 [4] Zolotov, M. Y. et al. (1997) Icarus, 130, 475-494. [5] Zolotov, M. Yu. et al., (1999) Planet. Space Sci. 47, 245-260 [6] Lewis, J. S. (1968) Icarus 8, 434–456 [7] Mueller, R. F. (1968) *Nature* 220, 55–57 [8] https://www1.grc.nasa.gov/space/geer/ [9] Gilmore, M. S., and Santos, A. R. (2023) *LPSC LIV*.

Acknowledgments

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